



**THE ANIMA PROBLEM – POSSIBLE LOCATION
OF THE THREAT LOCUS**

E. M. WIENZ & A. R. BORDON
LIFE PHYSICS GROUP - CALIFORNIA

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The problem addressed in this essay was originally encountered in the transcripts of four interviews between a bona fide journalist from the San Diego, CA., area who answered to the name of Sara and a self-described linguist of exotic languaging systems attached to the Labyrinth Group of the Advanced Contact Intelligence Organization (ACIO). The ACIO is an unacknowledged department of the National Security Agency (NSA). The linguist's name was Jaime Neruda. These interviews were part of a considerably larger body of work under the rubric WingMakers, which is also the name of the source website – <http://www.wingmakers.com>.

The main purpose of this essay is to address the problem of the location of the locus of threat to humankind said to be represented by the Anima. In one of his interviews with Sara, Neruda describes the Anima (which he describes by the Latin singular, Animus) as ontocyberenergetic life forms – that is, entities that are best understood as ontocybernetic organisms. He also offers one other clue: they hail from a planet in a solar system of a galaxy located 38 million light years from Earth. In this essay, we propose to develop a series of “candidate galaxies” as possible physical locus of the threat, and suggest on the basis of astrophysical and astronomical evidence the most likely candidate by name and location.

A SHORT INTERPRETATION OF THE PROBLEM

In the idiomaterial Unum in which the Earth human overfunction exists, spacetime as superdomain is, according to the Working Model, an extension-thought of the T-boundary, best described as a resonant harmonic thought of ultimate cause. In previous essays (Bordon 2004a, 2004b, 2004c 2004d, 2005a, 2005b), it was established that Nature as product of all holomovement in all superdomains caused by chains of causation with sui-gensis at the T-boundary can best be grasped as levels of manifestation (or LOMs).

Each LOM is identifiable to each of us human observers (as yet

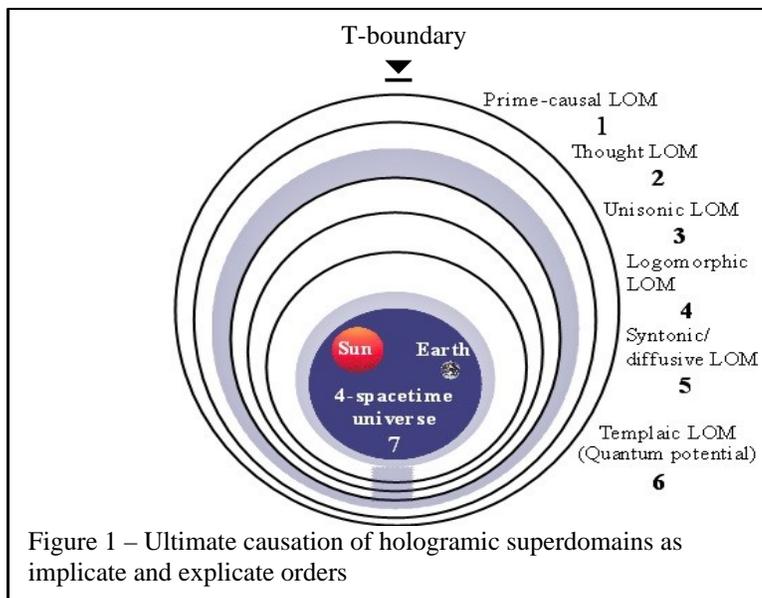


Figure 1 – Ultimate causation of hologramic superdomains as implicate and explicate orders

nonintegrated complex oscillating biominds) as ratios of time/space, each of which contains all possibilities.

Each time/space ratio as LOM holds all outcome probabilities in all levels of manifestation, thereby constituting literal histories; each LOM as time/space ratio contains time lines that include the five infinities (past timelike, future timelike, spacelike, past null, and future null infinities).

Just as motion is a hallmark in general and special relativity, holomovement is its equivalent in life physics: a holomovement is the motion-equivalent of any thought-essence undergoing transformation as relatively independent massless transform, any vector-intention expressed by any living relatively independent subtotaity (L-RIS), or any upward-causal or downward causal chain expressed by either the T-boundary or one or more L-RIS, or both. Holomovement, we have also come to discover recently, can be initiated by L-RISs in the inner three superdomains (Logomorphic, Syntonic-diffusive, and Quantum-potential). Life forms in 4-spacetime are referred to as ontobiological or ontocyberenergetic; all other life forms in superdomains higher than 4-spacetime are referred to simply as massless ontoenergetic entities (MOEs).

LEVELS OF MANIFESTATION (LOMs)

Each level of manifestation, as history (with a timeline) must contain all possibilities (as probable outcomes) of any and all holomovement in all superdomains. This is required by the Working Model for causality (as chains of causation) and causation to be a T-boundary metafunction to be so in any bundle of any aspect of any one or more superdomains.

A LOM is not a unit of space/time or time/space configured as a whole. Rather, it is a specifiable ratio of time/space in which idiomaterial manifestations, from biological cells to an entire universe, do take form. As media in which causal chains of thought-transforms and vector-intention wavefronts manifest, these affect the time/space ratio by the collapse of events in the timeline of the history stream. There is a direct relationship between observer and product or result of observation in a LOM. This relationship is referred to as *manifest product observership* (MPO) by the Working Model, which also treats it as result of observership in a LOM.

What then is the relationship between LOMs – the level of manifestation in which Earth exists and the level of manifestation in which the Animus planet is? The relationship is one of entangled vector-intentions. Within the confines of the framework disclosed by Neruda, both planets and their inhabitants know of one another. The entanglement of vector-intentions span a time period in both local times at the Earth LOM and at the Anima LOM. The analysis of vector-intention and discussion of the general threat as reported by Neruda is presented elsewhere (Bordon 2006).

In this essay, we will examine two sets of probable candidate environments in which the Anima LOM planet may be located. Using the distance given by Neruda, we will focus on galaxies in the Leo I and Leo II groups.

TARGET LOCATION CANDIDATE ENVIRONMENTS

FIRST CANDIDATE ENVIRONMENT. The first candidate environment is the Leo I group, composed of several galactic formation. Image 1 show three galaxies in the Leo I group. M95 is one of several galaxies within the M96 Group, a group of galaxies in the constellation Leo. The group also includes the Messier objects M95 and M96. M95 is one of several galaxies within the M96. Leo is one of

the most recognizable constellations in spring skies, although almost all of them are formed out of faint stars and have complicated shapes. A curve of stars arranged like a question mark turned over, "The Lion's sickle" represents Leo's head. One of famous meteor streams, Gamma Leonids, has a radiant point in the sickle, and the meteor shower was observed in November 1999. Alpha Leonis is named "Regulus" meaning of the little king, the star was well known in antiquity as Cor Leonis, the Lion's Heart. Leo is a neighbor of Coma Berenices, home of the Northern Galactic Pole, and so lies almost directly 'upward' through the plane of the Milky Way. Recent observations of planetary-nebula dynamics have given the intriguing result that M-105 (NGC 3379) is deficient in dark matter compared to most galaxies of its luminosity.

The prime candidate for target LOM in this group is M-105. It is the brightest elliptical galaxy in the Leo I or M96 group of galaxies and as such, approximately 38 million light years distant. It is of type E1, and often studied as a typical representative

of elliptical galaxies. Investigations with the Hubble Space Telescope of the central region of M105 have revealed that this galaxy contains a massive central object of about 50 million solar masses. Like many of the other galaxies in the Messier directory, M105 (aka NGC 3379) was observed with both Hubble Space Telescope, for a super-sharp view of the galaxy's center, and with ground-based telescopes, for a wider view of the regions farther from the galaxy's core.

Image 2 shows our second prime candidate – Messier object 96 (M-96). Spiral Galaxy M96 (NGC 3368), in Leo at a distance 38,000 (kly) from Earth, has a visual brightness 9.2 (mag). M96 is the brightest member of the Leo I group of galaxies, which is therefore also called the M96 group, and also includes M95, M105 as well as a number of fainter galaxies.

M96 is a type Sa spiral galaxy that is the brightest of the Leo 1 group of galaxies, which includes M95 and M105 as well as other fainter galaxies. It lies about 38 million light years away and it's bright inner portion extends about 66,000 light years in diameter. It has a fainter outer ring, which extend that diameter to about 100,000 light years. The inner disk is composed of an old population of yellow stars. The galaxy contains a significant amount of dust and blue knots of star forming regions.

Its distance was determined to be about 41 million light years (after corrections for the distance scale which are implied by the results of ESA's Hipparcos satellite) by Nial R. Tanvir (Tanvir et al. 1999) with the Hubble Space Telescope by



IMAGE 1 – Galaxies in the Leo I group (M-95, M96, and M-105)

Right Ascension:	10 : 46.8 (hours : minutes)
Declination:	+11 : 49 (degrees : minutes)
Distance:	38000 (* 1,000 light years)
Apparent Magnitude:	9.2
Apparent Diameter:	7.1 (arc minutes)

GALAXY M-105

observing Cepheid variables. Interpolated with the HST result of 35.5 million light years for its neighbor M95, we adopt a value of 38 million light years here for the whole group. At this distance, the apparent diameter of its brighter central region, 6 arc minutes, corresponds to a linear dimension of 66,000 light years. However, this galaxy has faint extensions, an outer ring of filaments (spiral arm fragments), which are connected to the bright visible part near the northwest end of the major axis. This ring has a diameter of at least about 9 arc minutes in the DSS image, corresponding to about 100,000 light years. The apparent visual brightness of 9.2 magnitudes corresponds to an absolute magnitude of -21.1.

Leo-I is the nearest group to Earth containing both bright spirals (M95 and M96) and a bright elliptical (M105) What appears to be a giant intergalactic HI ring orbits the central M105/NGC3384 galaxy pair and also appears to interact with M96. In a 22ksec ASCA SIS exposure of M96, Pedersen (2000) detected diffuse X-ray emission extending more than 10 arc minutes North of M96, in the direction of the HI ring. The morphology and spectral characteristics of the diffuse emission showed that M96 had recently interacted with the HI ring, indicating that M96, the HI ring and the central galaxy M105 are at the same distance from Earth (+/- 1.176 percent).). A first result was obtained for M95 in 1996-97 by the HST H0 Key Project Team (Graham *et.al.* 1997). They obtained a distance of 10.05+0.88 Mpc (32.8+-2.9 million light years), which after correction for the new Cepheid distance scale becomes 10.90+-0.95 Mpc (35.5+-3.1 million light years). As a mean value of the two distances obtained in the work of Tanvir and the H0 key project team, we will use a distance of 38 million light years here. From the distance derived, it is possible to gauge a diameter-velocity dispersion-relation for elliptical galaxies, namely by using the galaxies M105 and NGC 3377.

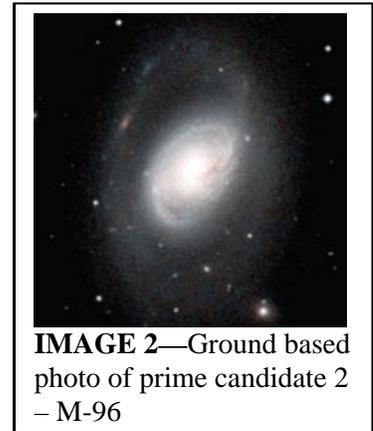


IMAGE 2—Ground based photo of prime candidate 2 – M-96

Further gnosive examination of the intergalactic environment of the Leo-I group showed the HI ring not to be an event-process that occurred naturally. At first decipherment, it appeared to be a stable and visible Einstein-Rosen bridges and the diffuse X-ray emission extending up to 17 arcminutes north of M96 was a byproduct of these bridges. It is visible to gnosive perception by a human neurosensor. Interestingly, the interaction between M96, M105 and NGC3384 was intermittent X-ray emission from the tube-like HI ring. Time-compressed observation in local time by a human neurosensors showed that the emissions occurred during periods of local time in which the ring became momentarily energized along its length from M96 to NGC3384, and then from NGC3384 to M105. The emissions were cyclical, meaning they would occur at roughly 60 minutes Earth time. The apparent cause at this stage of

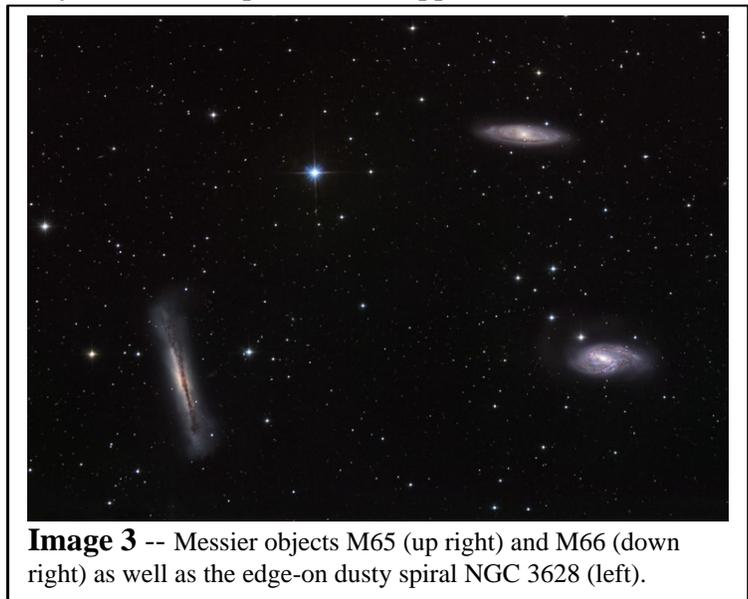


Image 3 -- Messier objects M65 (up right) and M66 (down right) as well as the edge-on dusty spiral NGC 3628 (left).

decipherment examination appeared to be an energetic bolus that would be hurdled in one direction.

SECOND CANDIDATE ENVIRONMENT. There is an additional set of candidates: the Leo Triplet: M-65, M-66, and NGC-3628. These three galaxies, the heart of the M66 group, form a beautiful

and photogenic group which is frequently referred to as the Leo Triplet; Halton Arp (1966) has included this group as number 317 in his Catalogue of Peculiar Galaxies (M66 is also Arp 16 and shows some nice H-alpha areas and a faint halo). Lying at a distance of some 35 million light years, it is probably physically related to the M96 group. See Image 3.

Our image of M65 was obtained with the Anglo Australian Telescope. Although it is close to and thus under the gravitational influence of its neighbors, M65 looks like a very "normal" Sa type spiral and seems to have felt little influence. It has a prominent central lens and tightly wound spiral arms, plus a prominent dust lane marking the facing edge. The luminous disk is dominated by a smooth old stellar population. Near the lane, some knots are visible, which, may be associated with star forming regions. The lane may hide regions

of star formation usually associated with such features in spiral galaxies.

M66 is considerably larger than its neighbor, M65, and has a well developed but not well defined central bulge, and is therefore classified

Sb. Obviously its spiral arms are deformed, probably because of the encounters with its neighbors. They seem to be distorted and displaced above the plane of the galaxy. Note how one of the spiral arms seems to

pass over the left side of the central bulge. Much dust is visible here, as well as a few pink nebulae, signs of star formation, near the end of one of the arms. Spiral galaxy M65 is a normal spiral galaxy not unlike our own Milky Way. In fact, M65 is a typical spiral galaxy of a type that could be found anywhere in the local universe. Given a morphological type of "Sa", M65 shows tightly wrapped spiral arms and a large nuclear central bulge. The central bulge stars are older and redder than disk stars, which appear more blue. Stars in the bulge of our own Milky Way Galaxy are also typically older and redder than stars in the disk where our Sun resides. Although it appears that M65's gravity has distorted M66's symmetry, M65's symmetry seems unaffected by M66. M65 is located roughly 35 million light years away. The galaxy is low in dust and gas, and there is little star formation in it; there has been some relatively recently in the arms. The ratio of old stars to new stars is correspondingly quite high. In most wavelengths it is quite uninteresting, though there is a radio source visible in the NVSS, offset from the core by about two arc-minutes. To the eye, M65's disk appears slightly warped, and its relatively recent burst of star formation is also suggestive of some external disturbance. Studies suggests that the two other galaxies in the Leo Triplet interacted with each other about 800 million years ago. Recent research by Zhiyu Duan suggests that M65 may also have interacted, though much less strongly. He also notes that M65 may have a central bar -- it is difficult to tell because the galaxy is seen from an oblique angle -- a feature which is suggestive of tidal disruption.

Table 1 – Redshift data for Messier galaxies

Galaxy/NGC	RA	Dec	Dist	RV_hel	RV_gal	
M65	3623	11:18:55.2	+13:05:35	35	807 +/- 3	723
M66	3627	11:20:15.0	+12:59:30	35	727 +/- 3	643
M96	3368	10:46:45.6	+11:49:18	38	897 +/- 4	797
M105	3379	10:47:49.5	+12:34:57	38	920 +/-10	823

Galaxy: Identification of Galaxy

RA, Dec: Right Ascension (in hours:minutes:seconds) and Declination (in degrees:minutes:seconds) for epoch J 2000.0

Dist: Distance in millions of light years (Mega light years, Mly)

RV_hel, RV_gal: Radial velocity in km/sec, **hel** is heliocentric and **wrt** the Galactic Standard of Rest (i.e., a position at rest wrt the Galactic Center)

NGC 3628 is the faintest and most difficult in the group, just faint enough to have escaped Messier's small telescopes (although it may be that his later instruments *might* have shown it, if he had ever looked at this place under very good conditions). NGC 3628 is seen edge-on. A conspicuous band of dark dust clouds form a broad equatorial band, which obscures the galaxy's bright central region, and hides most of the bright young stars in its spiral arms. The dust band, or belt, is obviously distorted and deformed in the outer regions of the galaxy. The reason for this deformation is evidently the gravitational interaction with its two bright neighbors, M65 and M66.

LEO TRIPLET'S LOM CONSIDERATIONS. NGC 3628, M65 (NGC 3623) and M66 (NGC 3627) are a well-known conspicuous grouping of galaxies in the constellation of Leo. The tight grouping are all within a 1 degree field in the sky and are physically and dynamically related as all three reside at the same distance of about 30 million light years. The distance between M65 and M66 is only 160,000 light years, the same distance between the Milky Way and its interacting companion, the Large Magellanic Cloud. The distance between NGC 3628 and the other two members of the triplet is slightly more than 300,000 light years. The galaxies exhibit evidence of mutual disturbance suggesting that the

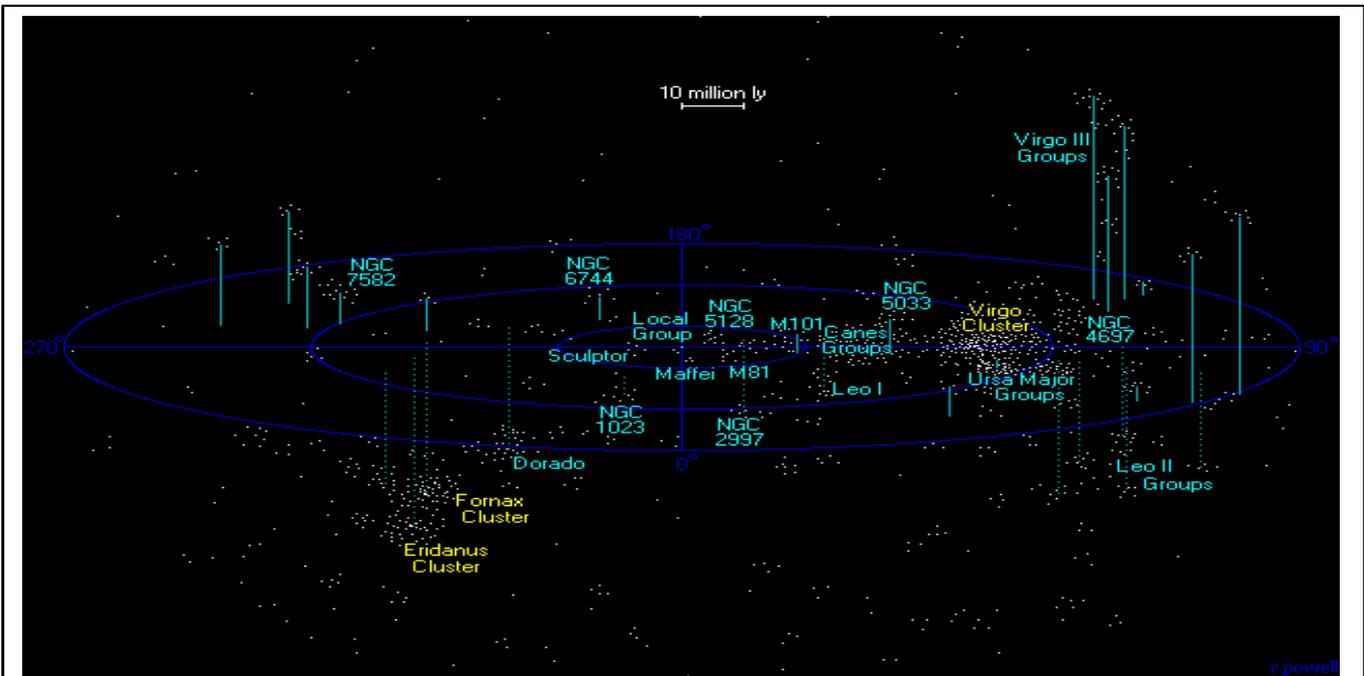


Image 4 – Relative positions of Leo I and Leo II galaxies relative to the Local Group

Leo triplet is a dynamically bound interacting system of galaxies.

The two galaxies which show the greatest evidence for a previous violent encounter are NGC 3628 and M66. Distinct morphological evidence exists that are believed to be the telltale signs of this interaction. NGC 3628 is noted for its faint optical plume that extends over 300 light years from its eastern tip and which contains some 500 million solar masses (about 15% of the total mass of NGC 3628). Star formation has been detected within the plume coinciding with the time of the encounter and therefore confirming its tidal origin. Surveys of neutral hydrogen gas using the tracer CO revealed a large bridge of gas between M66 and NGC 3628 as well as a gaseous plume coincident with the eastern optical plume observed in NGC 3628. Simulations have modeled the orbits of M66 and NGC 3628 and concluded that these galaxies came within 80,000 light years of each other some 800 million years ago,

and that encounter was responsible for the plumes and bridges we see now. The famous Leo triplet of Messier 65, 66 and NGC 3628 also demonstrates gravitational effects, especially on NGC 3628, which has a warped disk and a faint 260,000 light-year-long tail that is detected on the deepest photographs and at radio wavelengths. In addition the rotation curve for the eastern part of NGC 3628 falls sharply after reaching a maximum, which is believed to be a strong signature of a previous interaction as rotation curves typically remain flat or rise in normal spirals.

M66 also shows some disturbed features believed to be related to the remote encounter. A fossil starburst region near the nucleus of M66 dates to the same time as the proposed interaction slightly less than one billion years ago. The outermost southern gas clouds of the galaxy show a peculiar counterrotation in respect to the inner gas. Even more striking is an enormous gas concentration of about 400 million solar masses within the center of M66. Most likely this gas was accreted at the time of the closest approach of the two galaxies. There are several other asymmetries and distortions observed in the disks of NGC 3628 and M66 at both optical and radio wavelengths which are likely attributed to the previous encounter. Since multiple galactic rotations have occurred since the encounter it is difficult to trace backward the spatial relationships of the galaxies that existed at the time of the encounter.

In comparison to the other members of the Leo Triplet, M65 appears much more quiescent and stable. There are no obvious asymmetries, plumes or starbursts. One scenario raises the possibility that M66 and NGC 3628 existed as an orbiting bound system prior to the interaction with M66. M66 may have entered by way of "infall" into the small bound system of galaxies during the time period of the last encounter.

CANDIDACY CONSIDERATIONS

When the authors completed the first pass-through the cumulus on the Anima, their location, and other significant data, placement of the information derived from the cumulus by decipherment required the conduct of a search for candidates. By using information from a survey of the content of ancient texts recently conducted by the authors, the Leo constellation appeared to be highly significant in several texts from widely disparate ancient civilizations. Another important piece of information came from the data supplied by Neruda to Sarah, the linguist's choice of reporting channel. This concerned the distance between Earth and the galaxy where the Anima resided.

A survey of Leo I and Leo II galaxies under the Messier system immediately indicated galaxies M105 and NGC 3377 as prime candidates: both are 38 million light-years from Earth. However, since the authors already had the Anima signature from the first cumulus, it would be prudent to revisit each of the main Leo I objects, including the Leo Triplets. This made a total of seven galaxies: M105, NGC3377, M96, M95, M66, M65, and NGC3628. The authors then came across reports from the January 2006 American Astronomical Society's 207th Meeting held in Washington, D.C., held at the Marriott Wardman Park Hotel on Woodley Road.

A group of scientists from Cornell, the Kiev astronomical observatory, Lafayette College, the Wise Observatory, and the Arecibo Observatory, presented a paper quite relevant to the authors' work being reported in this essay. In the course of the presentation, it became evident that this report held vital information on the objectives of the authors' study. The group had looked at the Leo region because it offers a detailed view of several nearby groups of galaxies including Leo I at 10.4 Mpc and another slightly more distant structure within the Local Supercluster (Leo II). Leo I was of particular interest to both the scientific group and the authors because it contains both a large ring of intergalactic gas of

unknown origin (the Leo Ring) as well as a long tidal stream of stars and gas in the Leo Triplet. A catalog of HI line detections in the Leo region ($10^{\text{h}}30^{\text{m}} < \text{RA} < 11^{\text{h}}30^{\text{m}}$ and $+8^{\circ} < \text{dec} < +16^{\circ}$) had been made using data from the blind HI survey ALFALFA which takes advantage of the new, multi-feed array on the 305-m Arecibo telescope.

According to Stierwalt et al (2006), the ALFALFA detections were cross-referenced with optical catalogs and combined with targeted, single-pixel Arecibo observations of galaxies in the same area of the sky. The more sensitive observations targeted possible Leo Group candidates noted by Karachentsev and Karachentseva (members of the research team) through a visual inspection of POSS-II/ESO/SERC plates. Twenty-one of a possible thirty-five dwarf galaxy candidates were detected in HI including five background sources. Structure of the group and possible trends between characteristics like morphological class, HI mass, and distance from the group center were discussed in the course of the presentation. Of significance to the author, combining the ALFALFA data with the targeted, single-pixel observations allows for a direct comparison of optically and HI selected samples. This research, in combination with other work done by others (Zhiyu 2006, Pedersen 2000, Tanvir et al. 1999, Graham et al. 1997, Arp 1966), made possible the ultimate identification of the targeted LOM. The image above gives a sense of relative collocations of LOMs. See Image 4.

Therefore, as of this writing, the authors have increased the confidence level in the gnosive data on the M-105/NGC3377 by a factor of 3. Additional gnosive investigations conducted earlier this year have also confirmed signature considerations applicable to all objects in the Leo I group. To us, this indicates Anima presence in each of these galaxies. Gnosive examinations of M96, M95, M66, M65, and NGC3628 continue to date. No further data is yet available.

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